

# On the causes of Hard X-ray asymmetry in Solar Flares

R. Falewicz<sup>1</sup> and M. Siarkowski<sup>2</sup>

<sup>1</sup> Astronomical Institute of Wrocław University, 51-622 Wrocław, ul. Kopernika 11, Poland  
e-mail: falewicz@astro.uni.wroc.pl

<sup>2</sup> Space Research Centre, Polish Academy of Sciences, 51-622 Wrocław, ul. Kopernika 11, Poland  
e-mail: ms@cbk.pan.wroc.pl

Received xx xx, 2006; accepted xx xx, 2006

## ABSTRACT

*Aims.* Hard X-ray (HXR) emission in the footpoints of flaring loops often indicates an asymmetry, where brighter X-ray flux usually appears in the footpoint with a weaker magnetic field. This is explained by the fact that more electrons can reach the chromosphere in the weaker magnetic field. However, there are numerous exceptions to above rule, where the stronger HXR source is located in the loop's foot that is rooted in the stronger magnetic field. In our paper we analyse three such exceptional events from the paper by Goff et al. (2004). For all three events, we found evidence of the magnetic loops interactions near a brighter footpoint. Such loop interactions suggest that for these flares the energy release is located closer to one end of the loop.

*Methods.* We analysed the process of electron transport in detail in the dense flaring loop as a possible cause of the observed asymmetry. Using the flare parameters derived from observations, we calculated how the collisional energy losses in the loop reduce the number of electrons impacting the chromosphere at the other footpoint.

*Results.* We show that these calculations reproduce the observed flux ratios within the error limits in both footpoints for each of the analysed events. Thus we conclude that the injection of electrons near one of the footpoints, together with enhanced density in the flaring loop, can explain observed HXR footpoints asymmetry.

**Key words.** Sun: chromosphere – Sun: corona – Sun: flares – Sun: magnetic fields – Sun: hard X-rays

## 1. Introduction

Observation of hard X-ray (HXR) emission in solar flares is the main key to understanding the acceleration mechanisms of particles, localization of the site of energy release, and transport effects of the non-thermal electrons. Electrons accelerated anywhere in the corona move along magnetic field lines to the chromosphere where they deposit their energy. Part of the electrons' energy is converted via the bremsstrahlung process to HXR emission. As a result, during the impulsive phase of a solar flare, double (or multiple) HXR sources are often observed at the footpoints of flaring loops, an image now widely established from the Yohkoh and RHESSI observations.

Observed HXR footpoint fluxes often show asymmetry where the brighter source is usually located in a region of the weaker photospheric magnetic field (Sakao 1994, Kundu et al. 1995, Li et al. 1997, Aschwanden et al. 1999). This asymmetry is commonly interpreted as an effect of precipitation of the non-thermal electrons in a converging magnetic field (magnetic mirror effect). In the loop's leg with a weaker magnetic field (less effective magnetic mirror), more electrons are able to reach the chromosphere and vice versa. However, there

are numerous flares that do not fit the above scenario. One of five flares analysed by Sakao (1994) has a brighter footpoint located in the region of the stronger photospheric magnetic field. Asai et al. (2002) also report an example of a flare for which the brighter footpoint was found in a stronger magnetic field. Additionally, there are flares in which asymmetry changes with time. Examples of such events can be found in papers by Alexander & Metcalf (2001), Siarkowski & Falewicz (2004), and Falewicz et al. (2006).

Goff et al. (2004) studied the magnetic field strengths and hard X-ray brightness of 32 flares. They found that the brighter HXR footpoints coincide with the regions of stronger magnetic fields in 11 events. The authors designated these flares as N-type (from non-Sakao type), in contrast to S-type phenomena (for Sakao type, where a brighter footpoint is located in the region of the weaker photospheric magnetic field). According to the authors, the possible explanation of N-type events may be an asymmetry in the location of the acceleration site. If the acceleration site is located closer to the brighter footpoint, it may reduce the effects of magnetic field convergence, allowing more precipitation in this region than if the acceleration site was located at the loop apex.

Another factor that can be important for electron transport is matter density in the loop, as recently indicated e.g. by Veronig and Brown (2004). These authors report two flares observed with RHESSI for which the HXR emission comes mainly from the coronal part of the loops that are so dense as to be thick-target for electron energies up to 50 keV.

In this paper we explore the effect of the loop's density on the electron transfer in addition to an asymmetry of the injection place as the possible explanation of the Non-Sakao events. We analyse in detail three of the N-type flares studied by Goff et al. (2004) in which we have found evidence of loop interactions near the brighter of the HXR footpoints. This indicates that in these events a possible electron acceleration site may be located near one of the footpoints at the place where two magnetic loops interact. Using the observed HXR flux from an analysed flare, we calculate how the collisional energy losses reduce the number of electrons impacting the chromosphere at the other footpoint. We show that these calculations reproduce the observed flux ratios well in both footpoints for each of the analysed events.

In Sect. 2 we describe the analysed observations. The method of asymmetry modeling and calculations are presented in Sect. 3. In Sect. 4 we present the physical parameters of analysed flares. The results are discussed and we conclude in Sect. 5 and 6.

## 2. OBSERVATIONS

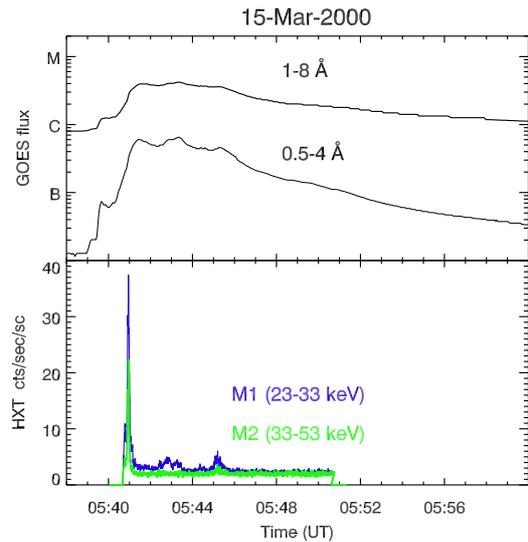
In this work we used data taken with the SXT grazing-incidence imaging telescope (Tsuneta et al. 1991), the HXT Fourier synthesis telescope (Kosugi et al., 1991), and *GOES* broadband X-ray fluxes measured in 1–8 Å and 0.5–4 Å wavelength bands.

The selection of events was based on N-type flares studied by Goff et al. (2004). From Goff's list we chose three events for which the configuration of the loops observed in SXT images indicated their possible interactions close to the brighter HXR footpoints. For all three analysed events, the count rates in the L (14–23 keV), M1 (23–33 keV), and M2 (33–55 keV) channels of HXT were well above the background level during the impulsive phase. Hard X-ray images of the flares were reconstructed using the standard Pixion reconstruction procedure (Metcalf et al. 1996). We used variable accumulation times assuming a threshold count rate of 200 counts in the M2 band (33–53 keV). The main advantage of the Pixion method is that it gives superior noise suppression and accurate photometry. In one case of a weaker event, the maximum entropy method (MEM) (Gull & Daniell 1978) was used to reconstruct the image. The MEM method also works for data with small photon statistics where the Pixion method failed.

The events are described below in detail:

### 2.1. 15 March 2000

This event occurred on 15 March 2000 around 05:40 UT. It was a C4.0 *GOES* class solar flare, which started at 05:40:30 UT, reached the maximum at 05:43 UT, and ended at 06:10 UT. This flare was located near solar disc centre (W15S17).



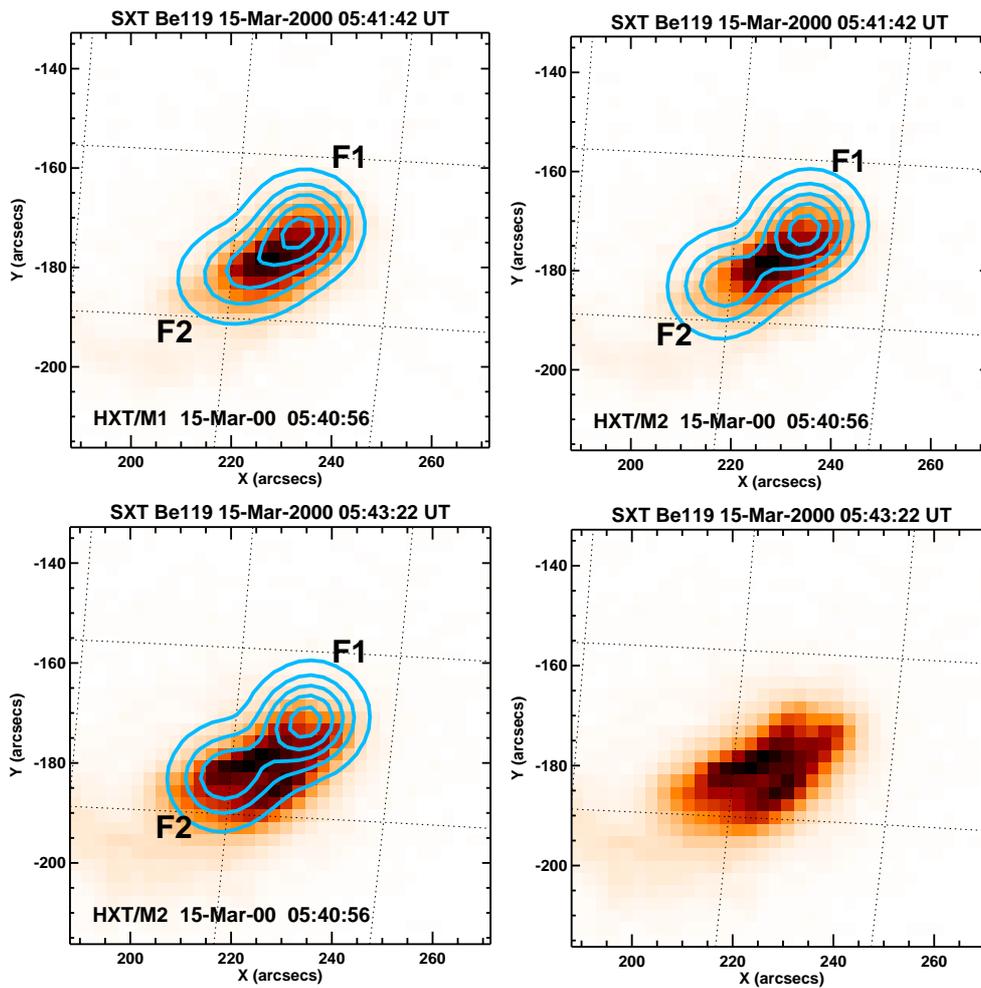
**Fig. 1.** X-ray light curves for the 15 Mar 2000 flare. *Top*: from *GOES*, *bottom*: from HXT M1 and M2 channels.

At the top of Fig. 1, *GOES* broadband X-ray fluxes are presented for this event measured in 1–8 Å and 0.5–4 Å wavelength bands. At the bottom of the same Figure, we show hard X-ray light curves in HXT channels M1 and M2. In Fig. 2 we present SXT images at two different times around flare maximum. These images were overlaid with intensity contours in M1 and M2 channels obtained at the peak of the HXR emission. The SXT images show a small flaring loop, whereas HXR emission tends to be concentrated at each end (denoted as F1 and F2 in Fig. 2) of this loop. This concentration is more evident at higher energies (M2).

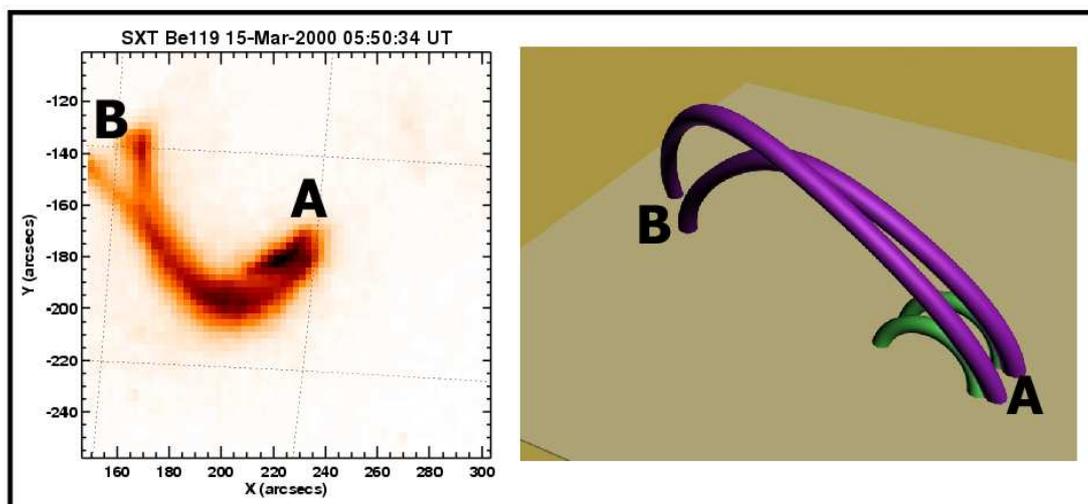
At the maximum of the HXR event ( $\sim 05:42$  UT), one small loop was observed in SXR (see Fig. 2 - two upper images). As the flare proceeded, a second parallel loop appeared at about 05:43 UT (see Fig. 2 - two lower images). As the presence of these two loops can be masked by HXT contours, we show the same SXT image in the lower panel of Fig. 2 with and without the HXT isocontours overlaid. The appearance of this second parallel loop is connected with a small brightening on both the SXR and HXR light curves in Fig. 1. Hard X-ray sources F1 and F2 were associated with the first SXR loop, which was the main flare loop. At lower HXR energies (channel L), emission comes generally from the F1 footpoint. Going to the higher energies (M1 and M2), the emission of the footpoints F1 and F2 dominates.

Later on during the flare decay, at least two weaker, extended loops appeared. This can be seen in Fig. 3 (left) where the composition of SXT images shows mutual locations of all loops during the flare. These extended loops lie parallel to two small loops, and the west region of their anchorage is located close to the footpoints of small loops (region A in Fig. 3 - left panel). This suggests I or Y-type interaction of the loops and we can assume that energy release in this event occurred close to the F1 footpoint of the flaring loop. This interpretation is justified by the fact that these larger loops are seen well before the flare (the last from 05:00 UT).

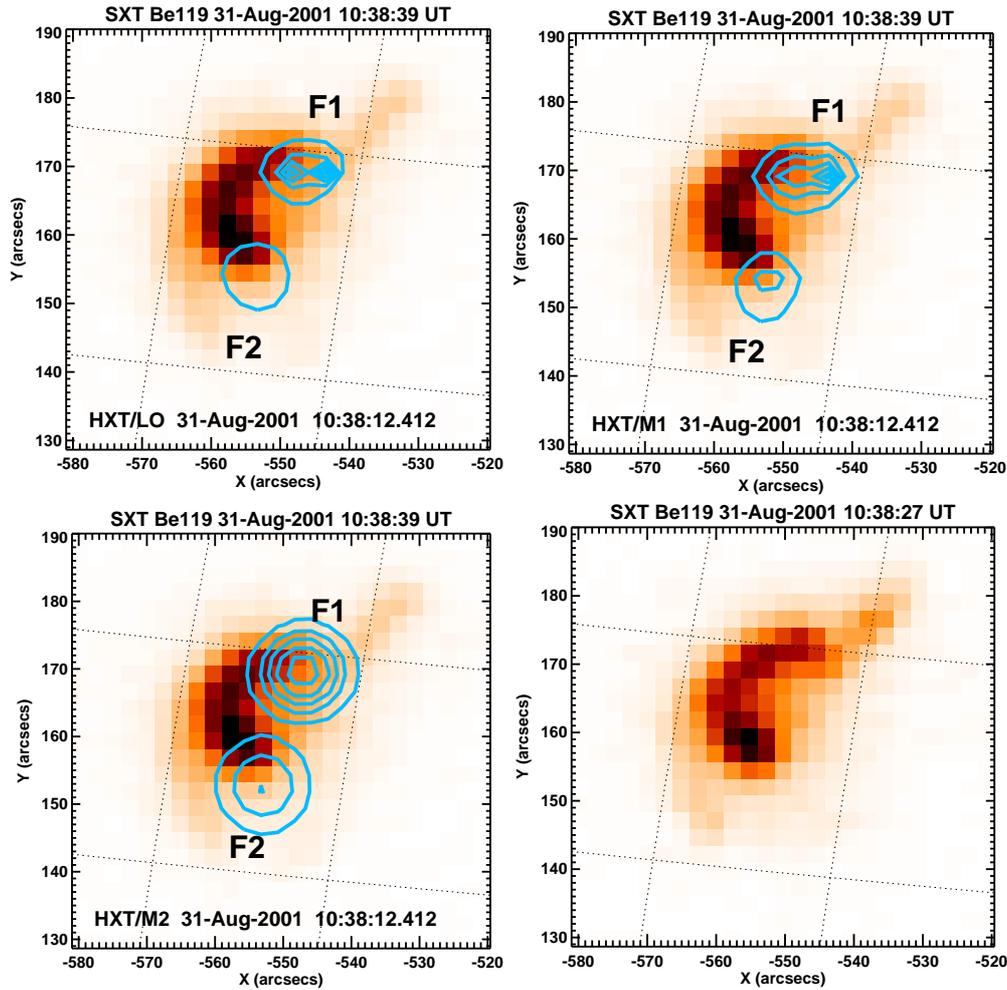
Schematic 3D models of the loop interaction in this flare is presented in the right panel of Fig. 3.



**Fig. 2.** (*Upper panel*): SXT soft X-ray image of the 15 March 2000 flare in the Be119 filter obtained at 05:41:42 UT. Contour levels at 0.1, 0.3,... 0.9 of the maximum flux show hard X-ray images at two HXT energy ranges M1 and M2. (*Lower panel*): SXT/Be119 image at 05:43:22 UT with and without the HXT/M2 contours overlaid.



**Fig. 3.** *Left*: Composition of SXT images showing mutual locations of the loops during the 15 March 2000 flare. *Right*: Schematic 3D models of the loop interaction in this flare. This picture was rotated and the point of view changed for clarity.



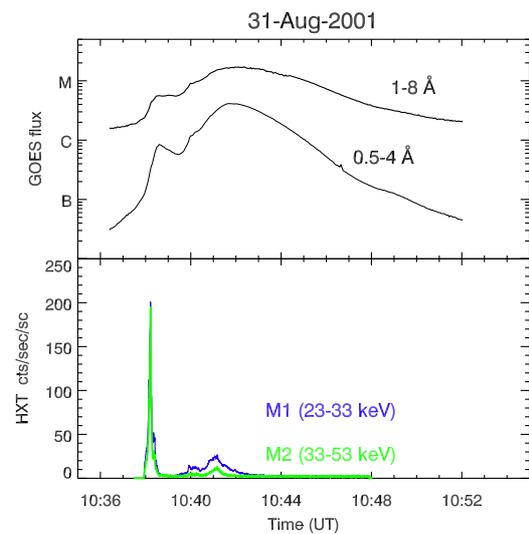
**Fig. 4.** Soft X-ray SXT/Be119 image of the 31 August 2001 at 10:38:39 UT overlaid with iso-contours of the hard X-ray emission observed at 10:38:12 UT at three HXT energy channels: L, M1 and M2. Contour levels are 0.1, 0.3,... 0.9 of the maximum flux. The bottom right-hand corner presents an SXT/Be119 image taken at 10:38:27 UT.

## 2.2. 31 August 2001

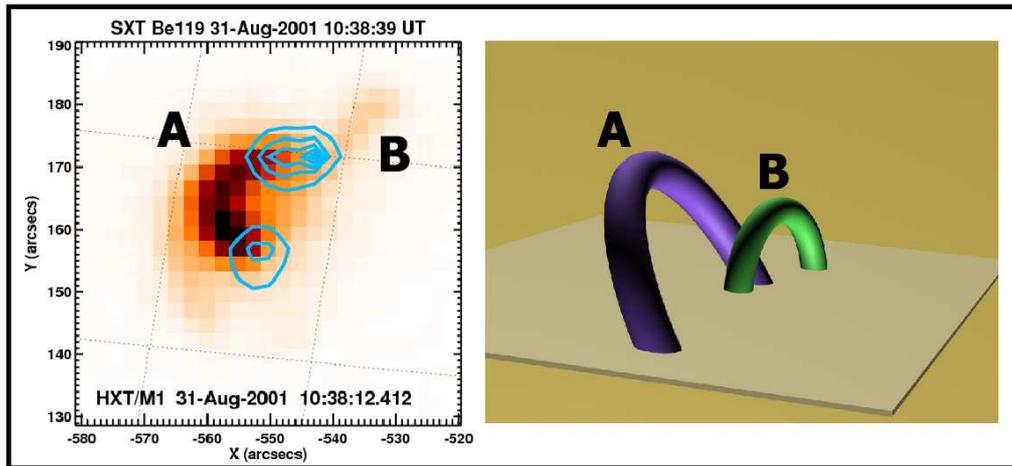
The flare on 31 August 2001 occurred at N15E37 and was an M1.6 *GOES* class event. It started at 10:35 UT, reached the maximum at 10:42 UT, and ended at 11:10 UT. Fig. 5 presents SXR and HXR light curves in channels M1 and M2.

In Fig. 4 we present the SXT image of this flare obtained during the rising phase at 10:38:39 UT with the Be119 filter. This image is overlaid with intensity contours obtained in LO, M1, and M2 images taken at about 10:38:12 UT. The visible sources of the hard X-ray emission can be recognised as footpoints of the magnetic loops (Falewicz et al. 2005). These footpoints denoted as F1 and F2 in Fig. 4 were connected by a loop visible in SXR.

Source F1 has two maxima in a common envelope of emission in the HXT channels LO and M1 and probably contains also a footpoint of the another weaker loop. The existence of this second loop was confirmed by the appearance of an extended emission, denoted as B in Fig. 6. In the bottom right-hand corner of Fig. 4, we present an SXT/Be119 image taken at 10:38:27 UT (12 seconds earlier than the three other).



**Fig. 5.** X-ray light curves for the 31 August 2001 flare. *Top*: from *GOES*, *bottom*: from HXT M1 and M2 channels.



**Fig. 6.** Left panel: Soft X-ray SXT image of the active region NOAA 9601 taken on 31 August 2001 at 10:38:39 UT. The main flaring loop, marked A, is accompanied by an extended emission, which gradually filled the shorter loop marked B. Right panel: The simple model of the interacting magnetic loops. The interaction of the loops occurred in "X" or "Y" configuration of the magnetic fields, low above the chromosphere. This picture was rotated and the point of view changed for clarity.

image the emission from the loop B is brighter and close to an F1 footprint. In fact a movie made of SXT images taken between 10:38:19 UT and 10:39:09 UT clearly reveals an ejection along a loop-like shaped structure that started exactly from source F1.

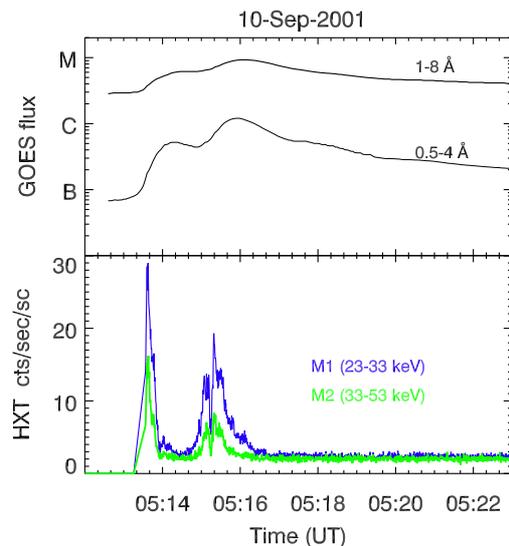
The observed structure of the HXR footprints, as well as loops seen in SXR, can be explained easily with a simple model of two interacting magnetic loops. The possible spatial configuration of the loops is presented in Fig. 6 (right panel). The region of interaction of the loops was located perhaps just above the chromosphere, in the closest vicinity of the F1 footprints (the configuration of the interacting loops was of either types X or Y), so for this flare we can assume again that the energy release took place low near the footprint F1.

### 2.3. 10 September 2001

This flare occurred on 10 September 2001 around 05:13 UT. It was a C5.8 *GOES* class solar flare, which started at 05:12 UT, reached the maximum at 05:16 UT, and ended at 05:23 UT. The flare was located near the solar disc centre (E14 S24).

*GOES* and HXT light curves of the event are presented in Fig. 7. The HXT light curves indicate that there are actually two energy release events in this flare, one lasting from 05:13:30 UT to 05:14:00 UT and the other from 05:14:50 UT to 05:16:20 UT. This is confirmed by the sequence of SXT images in which two nearly parallel small loops brightened one after the other during only a few minutes.

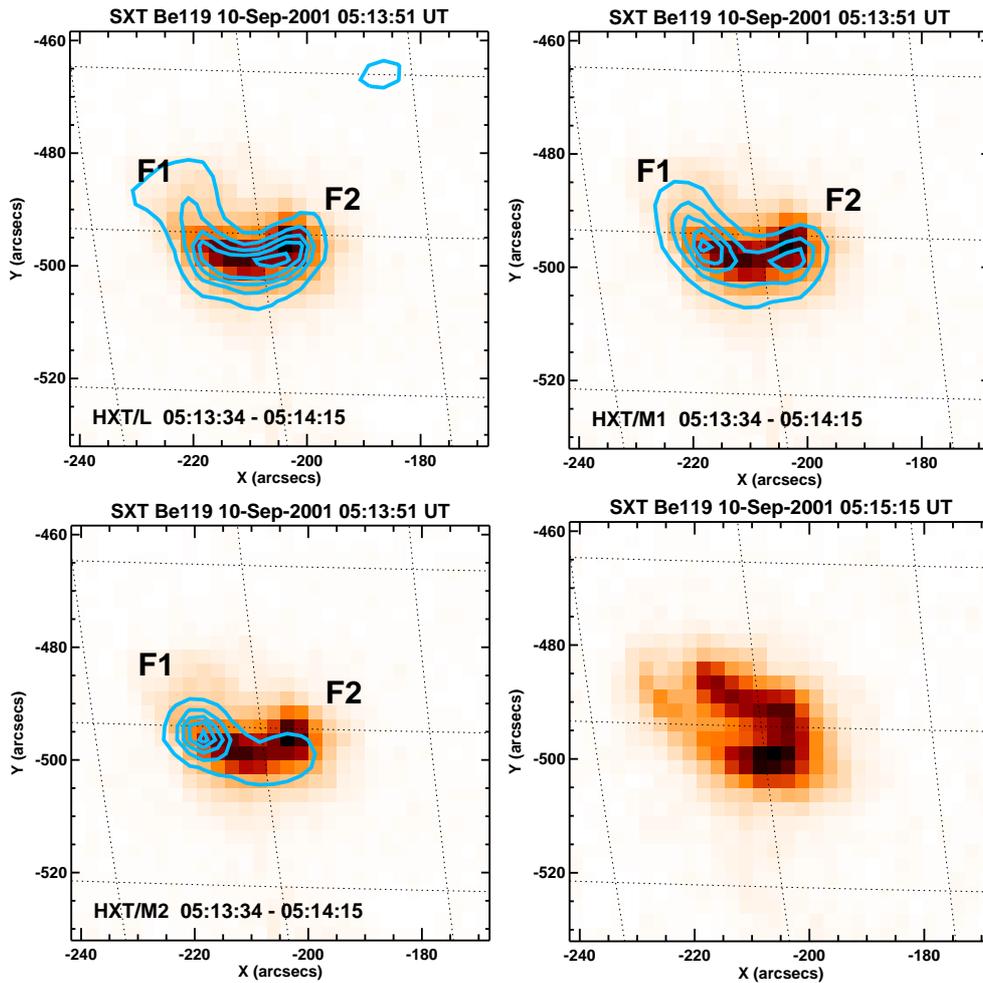
Fig. 8 shows an SXT soft X-ray image with overlaid hard X-ray contours at three HXT energy ranges: L, M1, and M2. At lower energies (channel L), emission comes from the whole loop. Going to higher energies (M1 and M2), emission at the footprints of this loop dominates. These footprints are again denoted by F1 and F2 in Fig. 8.



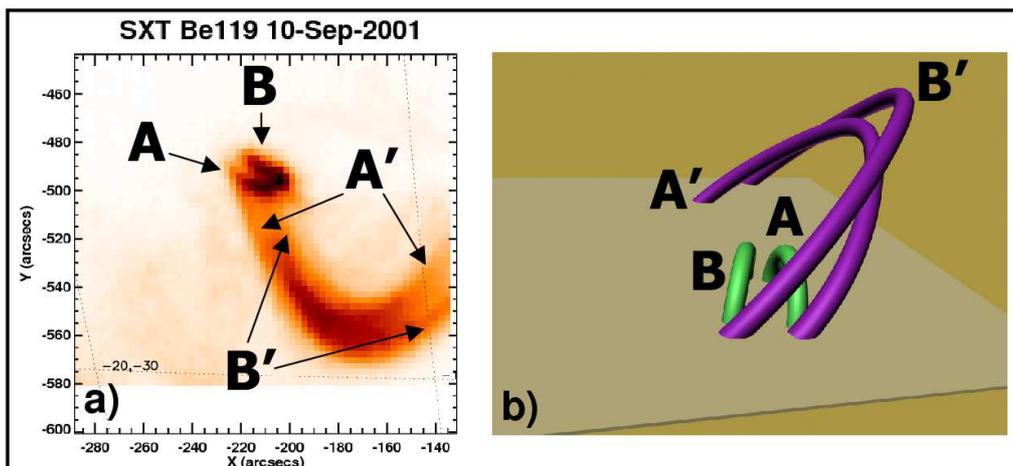
**Fig. 7.** X-ray light curves for the 10 September 2001 flare. *Top*: from *GOES*, *bottom*: from HXT M1 and M2 channels.

In the left panel of Fig. 9, compositions of SXT images is presented to show the mutual location of all loops during the event. The first flaring loop disappeared at about 05:15 UT (loop A), and at that time a second small parallel loop brightened (denoted as B) reaching soft X-ray flux maximum at 05:16 UT. When both small loops decreased in intensity another system of much larger loops appeared, in which two loops (denoted as A' and B') can finally be clearly distinguished. This system of loops is seen from the first SXT image, but they have a very low intensity compared to flaring loops A and B. They can also be recognised on the half and quarter resolution pre-flare images.

The presented image suggests the possible interaction of loop A with A' and B with B', near their F1 anchorages. Such interaction can cause energy release in the observed flare and can explain why the hard X-ray emission at the F2 footprints



**Fig. 8.** SXT soft X-ray images of 10 September 2001 flare at Be119 filter for two different times. Contour levels at 0.1, 0.3,... 0.9 of the maximum flux show hard X-ray images at three HXT energy ranges: L, M1, and M2. In the upper panel and the lower lefthand corner the SXT/Be119 image at 05:13:51 UT is presented when only one loop is visible. The bottom righthand corner presents an SXT/Be119 image taken at 05:15:15 UT, the time at which the first flaring loops decreased in intensity and the second one was reaching its maximum brightness.



**Fig. 9.** *Left:* Composition of SXT images showing mutual locations of the loops during the 10 September 2001 flare. *Right:* Schematic 3D models of the loops interaction for this flare. This picture was rotated and the point of view changed for clarity.

**Table 1.** Physical parameters of the analysed flares

Date		15-Mar-2000	31-Aug-2001	10-Sep-2001
	units			
<i>GOES</i> class		C4.0	M1.6	C8.7
maximum time	[UT]	5:43	10:42	5:16
Gamma, $\gamma$		3.68	2.94	4.1
Loop length, $L$	[ $10^9$ cm]	$2.6 \pm 0.2$	$2.0 \pm 0.2$	$2.0 \pm 0.2$
Footpoint area, $A$	[ $10^{17}$ cm <sup>2</sup> ]	$1.6 \pm 0.7$	$2.1 \pm 0.8$	$1.75 \pm 0.35$
Volume, $V$	[ $10^{26}$ cm <sup>3</sup> ]	$4.2 \pm 1.8$	$4.2 \pm 1.6$	$3.5 \pm 0.8$
Temperature ( <i>GOES</i> ), $T_e$	[MK]	$14.0^{+0.1}_{-0.5}$	$11.1^{+0.1}_{-0.1}$	$12.5^{+1.0}_{-1.1}$
Temperature (SXT), $T_e$	[MK]	$11.2^{+0.2}_{-0.2}$	$10.7^{+0.1}_{-0.1}$	–
Emission measure ( <i>GOES</i> ), $EM$	[ $10^{48}$ cm <sup>-3</sup> ]	$3.0^{+0.45}_{-0.77}$	$3.7^{+2.5}_{-0.73}$	$2.2^{+1.26}_{-1.49}$
Emission measure (SXT), $EM$	[ $10^{48}$ cm <sup>-3</sup> ]	$2.0^{+0.1}_{-0.1}$	$2.6^{+0.2}_{-0.2}$	–
Density ( <i>GOES</i> ), $n_e$	[ $10^{10}$ cm <sup>-3</sup> ]	$8.5^{+5.7}_{-6.0}$	$9.4^{+8.3}_{-6.1}$	$7.9^{+8.4}_{-6.7}$
Density (SXT), $n_e$	[ $10^{10}$ cm <sup>-3</sup> ]	$6.9^{+4.5}_{-4.5}$	$7.9^{+4.9}_{-4.9}$	–
Column depth ( <i>GOES</i> ), $N_L$	[ $10^{20}$ cm <sup>-2</sup> ]	$2.2^{+1.5}_{-1.6}$	$1.9^{+1.7}_{-1.2}$	$1.6^{+1.3}_{-1.3}$
Column depth (SXT), $N_L$	[ $10^{20}$ cm <sup>-2</sup> ]	$1.8^{+1.2}_{-1.2}$	$1.6^{+1.0}_{-1.0}$	–
Stopping energy ( <i>GOES</i> ), $E_{st}$	[keV]	$\approx 49.0$	$\approx 45.5$	$\approx 41.8$

of loops A and B is much weaker, if the loops are sufficiently dense to stop the non-thermal electrons on their way to these footpoints (Siarkowski & Falewicz 2005). Schematic 3D model of the loops interaction in the analysed event is presented in right panel of the Fig. 9.

### 3. Physical parameters of analysed flares

Table 1 lists the characteristics of the three events we selected. Time of maximum for each event in Table 1 is based on *GOES* flux in channel 1-8 Å. The spectral slope ( $\gamma$ ) was derived from the M1 and M2 HXT channel ratio. We calculated the loop length  $L$  using the distance of centres of gravity of HXT footpoints, assuming the semi-circular shape of the loop. When calculating  $L$ , we used the M1 channel for 10 September 2001 event and the M2 channel for two remaining flares. The footpoint area  $A$  was calculated as the circular cross-section of the loop seen on SXT images and volume as  $V = AL$ . We decided to use the definition of the footpoint area from SXT images because the Pixon method gives an accurate photometry of the sources, but the shape of reconstructed structures is often nearly circular. Additionally the diameters of those circles usually change with the source intensity (this can be seen for example in Fig. 4). The values of emission measure  $EM$ , temperature  $T$ , and electron density  $n_e$  were derived using 3-second *GOES* data (Thomas et al., 1985) and, for comparison, also from the SXT data for the time of HXT image during the impulsive phase. In only one case (10-Sep-2001), we could not derive temperature and emission measure from the SXT images because observations were unavailable at this time (impulsive phase). For uniformity of the analysed data, we used the *GOES* diagnostic that was available for all three events. Additionally *GOES* data can be easily interpolated to the time where HXT images was reconstructed. This is particularly important during the impulsive phase when the SXT flux increases quickly. However, if possible we also calculated  $n_e$  and  $N_L$  from SXT data for comparison. We calculated the column depth  $N_L$  using *GOES* and the SXT data as  $N_L = n_e L$ . The energy of the

electrons that are collisionally stopped by the column depth is  $E_{st} = (3 \times K \times N_L)^{1/2}$  (e.g. Fisher, Canfield, and McClymont 1985), where  $K = 3.64 \times 10^{-18}$  keV<sup>2</sup> cm<sup>2</sup>. The errors given in Table 1 were calculated using statistical uncertainties. Errors in geometrical parameters were calculated assuming 1 pixel (2.544 arc sec) uncertainties. Our errors are rather high, but it should be noted that we have accepted maximum errors as far as possible for assumed geometry.

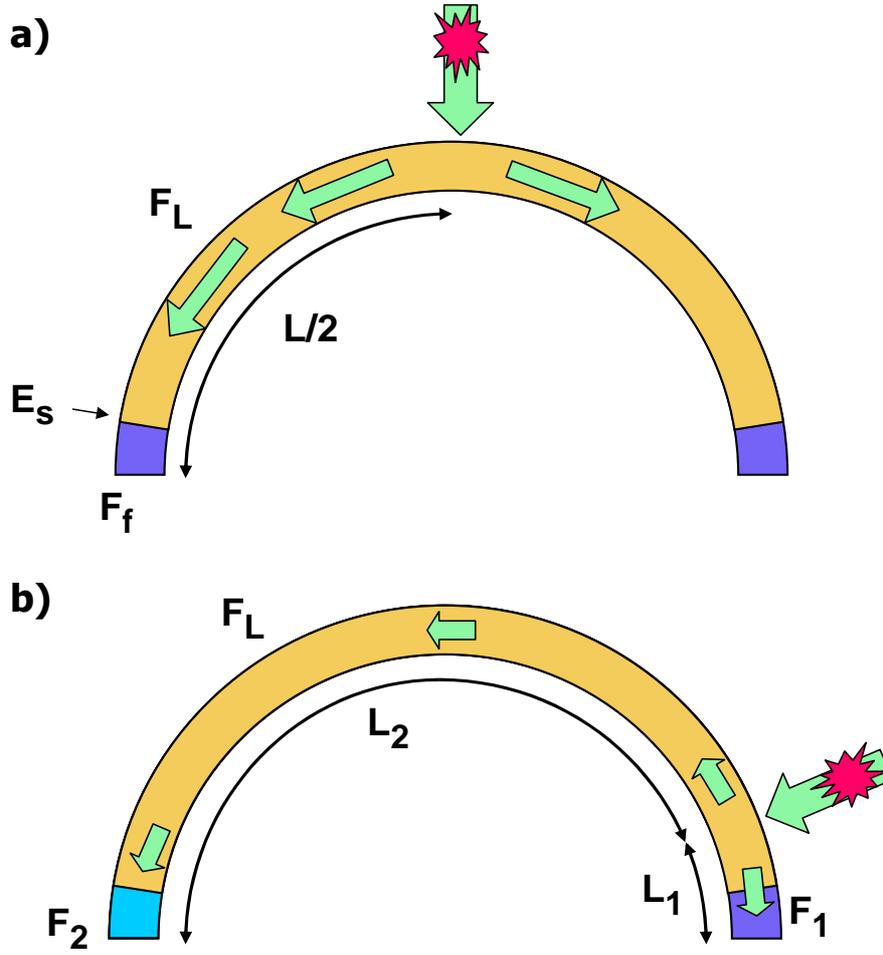
### 4. Theoretical interpretation

Usually the source of the energy release is located high in the corona and the site of electrons injection to the flaring loop lies near the loop top (Fig. 10a). If the loop is dense enough, then electrons collisionally lose a significant part of their energy and the emission at the footpoints  $F_f$  will not be pure thick-target emission, but will be reduced to:

$$F_f = F_{th} - F_L \quad (1)$$

where  $F_{th}$  is the thick-target flux and  $F_L$  emission from the loop's leg.

The effect of collisional energy loss can be especially important in the case where electrons are injected near one of the footpoints (Fig. 10b) as suggested by the observations described above. Then plasma filling the loop can reduce the number of electrons impacting the chromosphere at the other footpoint, causing the observed HXR flux asymmetry. This model is used in the rest of the paper. We assume that the electrons are injected in the same proportions in both the up and down directions along the loop and that the HXR flux (channels M1 and M2) in the brighter footpoint  $F_1$  is a thick-target emission, i.e.:  $F_{th} = F_1 = F_2 + F_L$  (see Eq. 1), where  $F_1$  is the flux in the brighter footpoint,  $F_L$  the flux in the whole loop legs (without footpoints; further shortly: loop), and  $F_2$  the flux at the other footpoint. The electrons travel from the place of injection to the footpoints  $F_1$  and  $F_2$  along the lengths  $L_1$  and  $L_2$ , respectively. Hence in our model  $L_1 + L_2 = L$ , but we assume that  $L_1 \ll L_2 \approx L$ . To calculate  $F_2$  we adopted here an



**Fig. 10.** Schematic model of the electron transfer in the flaring loop: a) in the case where energy release is located at or above the loop apex, and b) when energy release is located near one of the footpoints. For more details see text.

intermediate thin-thick target model of Wheatland & Melrose (1995) in which low-energy electrons are collisionally stopped by column depth  $N_L = n_e L$ . Using the notation of Wheatland & Melrose (1995), we can write the ratio of fluxes measured at both footpoints as

$$r = \frac{F_1}{F_2} = \frac{B(\delta - 1, \frac{1}{2})}{(\delta - 2) G(E_c/\epsilon, \delta)} \quad (2)$$

where:  $\epsilon$  is a photon energy,  $E_c = (2KN_L)^{1/2}$ ;  $K$  is a function of the Coulomb logarithm (assumed as constant);  $\delta$  is a power-law electron spectral index ( $\delta = \gamma + 1$ , where  $\gamma$  is observed photon spectral index);  $B(p, q)$  is the beta function; and

$$G(E_c/\epsilon, \delta) = \int_0^1 x^{\delta-3} (1 + (E_c/\epsilon)^2 x^2)^{-(\delta-1)/2} \ln \left[ \frac{1 + \sqrt{1-x}}{1 - \sqrt{1-x}} \right] dx \quad (3)$$

Using an observed photon spectral index  $\gamma$  and the resulting loop's parameters  $n_e$  and  $L$ , we can calculate the flux ratio and compare it to those obtained from HXR images. Due to the projection effect, the flux measured in footpoints  $F_1$  and  $F_2$  is not true footpoint emission but the sum of the footpoint emission and part of the emission from the loop's legs. As the

stopping energy  $E_{st}$  in each flaring loop is quite high (see Table 1), the emission from the legs is important in both HXT channels M1 and M2, so the calculation of the observed fluxes  $F_1$  and  $F_2$  values is more complicated. As the total flare emission is a thick-target, it follows from our model that  $F_1 = F_{total}/2$  and then  $F_2 = F_{total}/2 - F_L = F_1 - F_L$ . The value  $F_L$  was calculated by evaluating the intensity near the loop top and the whole loop volume. The values of the measured and calculated fluxes and their ratios for each of the flares are presented in Table 2. The range of the  $F_L$  values results from the errors in the cross section of the loop. We also calculated the stopping column depth for each energy channel, which is the column of matter in which electrons totally rest their energy due to the collisions - only electrons with energy  $E > E_{st}$  can penetrate deeper. We denote these quantities as  $N_{M1}$  and  $N_{M2}$  for channels M1 and M2 (effective energies 29 keV and 44 keV, respectively). In Table 2 we give the ratio  $q_{M1} = N_{M1}/N_L$  and  $q_{M2} = N_{M2}/N_L$  for each flare (*GOES* and *SXT* data). For example  $q_{M2} \approx 0.5$  means that a large fraction of electrons with energy  $\approx 44$  keV are collisionally stopped near loop's top and when  $q_{M2} \geq 1.0$  means that these electrons freely go to the other footpoint.

**Table 2.** Photometric and model parameters of the analysed flares

Date	15-Mar-2000			31-Aug-2001		10-Sep-2001	
<b>M1 channel</b>							
	UNITS	OBSERVATIONS	MODEL	OBSERVATIONS	MODEL	OBSERVATIONS	MODEL
$F_1$	<i>ph/sec</i>	79.2	–	37.8	–	100.4	–
$F_L$	<i>ph/sec</i>	–	44.1	3.9 – 9.8	8.02	52.7 – 81.6	74.12
$F_2$	<i>ph/sec</i>	–	–	28 – 33.9	–	19.4 – 47.7	–
$r_{M1} = F_1/F_2$		–	2.26	1.11 – 1.35	1.27	2.1 – 5.2	3.82
$q_{M1} = N_{M1}/N_L$ ( <i>GOES</i> )		$0.35^{+0.24}_{-0.25}$		$0.41^{+0.36}_{-0.26}$		$0.48^{+0.39}_{-0.39}$	
$q_{M1} = N_{M1}/N_L$ ( <i>SXT</i> )		$0.43^{+0.28}_{-0.28}$		$0.48^{+0.30}_{-0.30}$		–	
<b>M2 channel</b>							
	UNITS	OBSERVATIONS	MODEL	OBSERVATIONS	MODEL	OBSERVATIONS	MODEL
$F_1$	<i>ph/sec</i>	40.5	–	25.4	–	27.0	–
$F_L$	<i>ph/sec</i>	7.9 – 21.8	13.8	1.2 – 5.3	2.60	14.0 – 21.4	14.06
$F_2$	<i>ph/sec</i>	18.7 – 32.6	–	20.1 – 24.2	–	4.6 – 13.0	–
$r_{M2} = F_1/F_2$		1.24 – 2.16	1.52	1.05 – 1.26	1.11	2.1 – 5.9	2.08
$q_{M2} = N_{M2}/N_L$ ( <i>GOES</i> )		$0.80^{+0.55}_{-0.59}$		$0.93^{+0.83}_{-0.59}$		$1.11^{+0.90}_{-0.90}$	
$q_{M2} = N_{M2}/N_L$ ( <i>SXT</i> )		$0.99^{+0.65}_{-0.65}$		$1.11^{+0.69}_{-0.69}$		–	
Density ( <i>GOES</i> ), $n_e$	$10^{10} \text{ cm}^{-3}$	$8.5^{+5.7}_{-6.0}$	6.9	$9.4^{+8.3}_{-6.1}$	3.2	$7.9^{+8.4}_{-6.7}$	14.5
Density ( <i>SXT</i> ), $n_e$	$10^{10} \text{ cm}^{-3}$	$6.9^{+4.5}_{-4.5}$	6.9	$7.9^{+4.9}_{-4.9}$	3.2	–	14.5

## 5. Results

The interaction of magnetic structures and conversion of magnetic free energy into plasma heating and particle acceleration is commonly regarded as the origin of solar flares. The configurations of the interacting loops observed on the Sun may be described using three main models: the I-type coalescence (interaction of two nearly parallel loops), the Y-type coalescence interacting loops (in contact only along a part of their lengths), and the X-type coalescence when some parts of two loops crossed approximately perpendicularly (Sakai and de Jager 1996). Most analyses of the observational data concerning the coalescence of loops describe the I-type interactions (e.g., Inda-Koide et al. 1995; Takahashi et al. 1996), some are related to the Y-type (e.g., Farnik et al. 1996; Hanaoka 1994), and only a few describe X-type (e.g., de Jager et al. 1995; Falewicz & Rudawy 1999; Rudawy et al. 2001). All three flares may be interpreted as resulting from the interaction of magnetic loops, as described above. Localisation of this interaction near a brighter footpoint suggests that the energy release is located low, close to one end of the loop. Using physical parameters obtained from observations and our model of the energy release near one of the footpoint, we calculated the ratio of the flux emitted by footpoints F1 and F2 (see Table 2).

For the 15 March 2000 flare, we cannot determine the observed ratio ( $r_{M1}$ ) in channel M1, but in channel M2 the obtained value ( $r_{M2}$ ) of 1.52 corresponds well to the observed range 1.24 – 2.16 (see Table 2). The calculated value of ratio  $q_{M1} = 0.35$  for the M1 channel indicates that most of electrons can reach the top of the loop only. In fact, in Fig. 2 (left panel) we can see that the emission in the M1 channel comes generally from half of the loop. In channel M2 this ratio (close to  $\sim 1$ ) indicates that electrons can reach the other footpoint as can be seen also in Fig. 2 (centre and right panel). We have obtained a similar result for  $q_{M1}$  and  $q_{M2}$  for SXT data.

In the case of the 31 August 2001 event, the ratio calculated from the model of the fluxes emitted by the footpoints F1 and F2 is  $r_{M1} \approx 1.27$  in the M1 channel and  $r_{M2} \approx 1.11$  in the M2 channel. These values correspond very well to the observed range values 1.11 – 1.35 and 1.05 – 1.26 for the same channels (see Table 2). The calculated values of ratio  $q_{M1}$  for M1 channel both for *GOES* and *SXT* data indicated that most of the electrons can only reach the top of the loop (see Table 2). The calculated values  $q_{M2}$  (*GOES* and *SXT* data) indicate that in channel M2 the majority of electrons can reach the other footpoint.

During the 10 September 2001 flare, footpoint F2 in Fig. 8 is brighter in channel M1 than in M2. This may be seen as contradicting the density effect, but it should be possible for most softer spectra with a higher spectral index. In fact for this flare,  $\gamma = 4.1$ , the higher value of all three events. It should also be remembered that by looking at the uniform intensity loop from above we will observe more emission at the footpoints than in the top as a projection effect.

Using our model we obtained for M1 channel  $r_{M1} \approx 3.8$ . This last value corresponds well to the range 2.1–5.2 calculated from the M1 image. For the channel M2 the calculated ratio is  $r_{M2} \approx 2.1$  and still within the observed range 2.1 – 5.9.

To obtain values of the calculated ratio  $r = F_1/F_2$  in the observed range, we changed the density within the error limits so that we get a best fit to the observed parameter e.g.  $F_1$ ,  $F_L$ , and  $r$ . The adopted values of density for this and previous events are given in Table 2. As in the aforementioned cases we can observe very similar values of  $q$  (only for the *GOES* data because *SXT* data were unavailable). The calculated value of ratio  $q_{M1} = 0.48$  for the M1 channel indicated that most of the electrons can only reach the top of the loop. In channel M2, this ratio (of  $\sim 1$ ) indicates that most of the electrons can reach the other footpoint.

The values of  $n_e$ ,  $N_L$ ,  $q_{M1}$ , and  $q_{M2}$  obtained from both the *GOES* and *SXT* instruments are quite similar for the first two events (see Tables 1 and 2).

## 6. Discussion and conclusions

All three of the flares are so-called Non-Sakao types, i.e., the HXR emission is brighter in the footpoint with a stronger underlying magnetic field. According to Goff et al. (2004) these events can be explained by an asymmetry in the location of the acceleration site. If the acceleration site is located closer to the brighter footpoint it may reduce the effects of magnetic field lines convergence over the shorter distance, allowing more precipitation in this region than if the acceleration site was located at the loop apex. However, as pointed out by the authors, there is no simple relation between the HXR flux of any individual footpoint to the magnetic flux density beneath it, so that the magnetic field strength in the footpoints can only be one of the factors that can determine asymmetry scenario. It is possible, for example, that a level of a strong field convergence, which is important for the mirroring of electrons, lies below the level where HXR emission is formed. Then, other effects, e.g. density in the loop, should be responsible for the observed asymmetry.

In fact, the densities of the analysed flaring loops are quite high, well above  $10^{10}\text{cm}^{-3}$ . Lately Veronig and Brown (2004), based on *RHESSI* observations, reported a new class of solar flare hard X-ray sources in which the emission is mainly in a coronal loop that is so dense as to be collisionally thick at electron energies up to  $\approx 50$  keV. While the source of the energy released in these events is probably located near the loop top, both footpoints only have a weak emission.

In the analysed flares we observed evidence that the magnetic loops interact. These interactions take place low in the corona near one of the loops' anchorage, so we can assume that the region of energy release and the source of the non-thermal electrons are placed near one of the footpoints. These electrons produce observed HXR in both footpoints. However, a fraction of the electrons going to the other footpoint can be collisionally stopped in a dense loop leading to the observed asymmetry. We used the physical parameters of flaring loops obtained from *GOES*, *SXT*, and *HXT* observations to calculate how the density in the loop affects the number of nonthermal electrons reaching the footpoints.

Our calculations indicate that we can explain this observed asymmetry mainly as an effect of the enhanced density in the flaring loop and of the injection of electrons near one of the footpoints. We show that these calculations reproduce the observed flux ratios well in both footpoints for each of the events we analysed. However, it should be remembered that the errors of the density and loop cross-section are large, so our results are only rough estimates. Additionally, while the observed loops are small, the effects of projections are important, and a large part of the loop's emission affects the emission of the footpoints. Another topic for further discussion is that in our simple model we do not consider the changes in average pitch angle of the electrons (scattering) as they traverse the dense region. In the end, it is also very probable that there is a com-

ination of magnetic-mirror and density effects. This last point should be especially emphasised: our density errors are sufficiently large and the magnetic field effects sufficiently unexplored. This problem needs further observations and analysis.

*Acknowledgements.* We are grateful to the anonymous referee for useful comments and suggestions. This work was supported by the Polish Ministry of Science and High Education, grants N203 022 31/2991 and 1P03D 017 29.

## References

- Alexander, D. & Metcalf, T. R., 2002, *Solar Physics*, 210, 323  
 Asai, A., Masuda, S., Yokayama, T., et al. 2002, *ApJ*, 578, L91  
 Aschwanden, M.J., Fletcher, L., Sakao, T., Kosugi, T., Hudson, H., 1999, *ApJ*, 517, 977  
 Falewicz, R. & Rudawy, P., 1999, *A&A*, 344, 981  
 Falewicz, R., Rudawy, P., Siarkowski, M., 2005, *ESA SP-600*, 121  
 Falewicz, R., Siarkowski, M., and Berlicki, A., 2006, *Advances in Space Research*, in press  
 Farnik, F., Hudson, H., Watanabe, T., 1996, *Solar Physics* 165, 169  
 Fisher, G., Canfield, R., McClymont, A. 1985, *ApJ*, 289, 414  
 Goff, C.P., Matthews, S.A., van Driel-Gesztelyi, L., Harra, L.K., 2004, *A&A*, 423, 363  
 Gull, S. F. & Daniell, G. J., 1978, *Nature*, 272, 686  
 Hanaoka, Y., 1994, *ApJ*, 420, L37  
 Inda-Koide, M., Sakai, J.I., Koide, S.J., et al., 1995, *PASJ* 47, 323  
 de Jager, C., Inda-Koide, M., Koide, S.J., & Sakai, J.I., 1995, *Solar Phys.* 158, 391  
 Li, J., Metcalf, T. R., Canfield, R. C., Wlser, J. P. 1997, *ApJ*, 482, 490  
 Kosugi, T., Masuda, S., Makishima, K., Inda, M., Murakami, et al., 1991, *Solar Phys.*, 136, 17  
 Kundu, M. R., Nitta, N. White, S. M., Shibasaki, K., Enome, et al., 1995, *ApJ*, 454, 522  
 Metcalf, T. R., Thomas, R., Hudson, H. S., Hugh, S., Kosugi, T., et al., 1996, *ApJ*, 466, 585  
 Rudawy, P., Falewicz, R., Mandrini, C.H., Siarkowski, M., 2001, *A&A* 372, 1037  
 Sakai, J.I., & de Jager, C., 1996, *Space Science Rev.* 77, 1  
 Sakao, T., 1994, Ph.D. thesis, University of Tokyo  
 Sakao, T., Sakao, T., Kosugi, T., Masuda, S., Yaji, K., Inda-Koide, M., and Makishima, K., 1996, *Adv. Space Res.*, 17, 67.  
 Siarkowski, M. & Falewicz, R., 2004, *A&A*, 428, 219  
 Siarkowski, M. & Falewicz, R., 2005, *AIP Conf. Proceedings*, 801, 403  
 Takahashi, M., Watanabe, T., Sakai J., et al., 1996, *PASJ* 48, 857  
 Thomas, R.J., Starr, R., Cranell, C.J., 1985, *Solar Phys.* 95, 323  
 Tsuneta, S., Acton, L., Bruner, M., Lemen, J., Brown, W., et al., 1991, *Solar Phys.*, 136, 37  
 Veronig, A. M., Brown, J. C. 2004, *ApJ*, 603, L117  
 Wheatland, M. S. & Melrose, D. B., 1995, *Solar Physics*, 158, 283